1870MNRAS..30..183S

Note on Dr. Wolfers' Paper. Comparison between the Places of the Greenwich Seven-year Catalogue and Radcliffe Catalogue, 1860, *for the Stars in Table I.* By E. J. Stone, Esq.

	Greenwich Mean Right Ascension.			Rad.	G.—R.	Mean N.P.D.			Rad.	G.—R.
	h	m	s	″	″	°	′	″	″	″
α Andromedæ	0	1	9·44	9·38	+ 0·06	61	40	57·79	56·3	+ 1·5
γ Pegasi	0	6	1·79	1·77	+ 0·02	75	35	42·37	39·0	+ 3·4
α Cassiopeiæ	0	32	35·06	35·14	− 0·08	34	13	51·94	52·7	− 0·8
α Arietis	1	59	17·28	17·25	+ 0·03	67	12	5·60	3·6	+ 2·0
α Ceti	2	54	57·83	57·89	− 0·06	86	27	42·81	42·6	+ 0·2
α Persei	3	14	20·72	20·74	− 0·02	40	38	26·72	27·0	− 0·3
α Tauri	4	27	53·43	53·43	0·00	73	46	32·97	31·6	+ 1·4
α Aurigæ	5	6	21·12	21·06	+ 0·06	44	8	57·40	57·2	+ 0·2
β Orionis	5	7	48·63	48·62	+ 0·01	98	21	59·90	59·2	+ 0·7
β Tauri	5	17	26·67	26·69	− 0·02	61	30	54·08	51·8	+ 2·3
α Orionis	5	47	35·59	35·58	+ 0·01	82	37	21·95	21·2	+ 0·8
α Canis Majoris	6	38	58·66	58·63	+ 0·03	106	31	36·23	36·6	− 0·4
α Geminorum	7	25	39·31 39·73	39·29 39·74	+ 0·02 − 0·01	57	48	33·16 30·65	31·8 29·5	+ 1·4 + 1·2
α Canis Minoris	7	31	58·34	58·35	− 0·01	84	25	9·79	9·2	+ 0·6
β Geminorum	7	36	44·68	44·64	+ 0·04	61	38	21·25	20·3	+ 1·0
α Hydræ	9	20	42·44	42·43	+ 0·01	98	3	13·76	15·0	− 1·2
α Leonis	10	0	54·78	54·71	+ 0·07	77	21	0·62	59·4	+ 1·2
α Ursæ Majoris	10	55	3·40	3·51	− 0·11	27	29	39·12	38·6	+ 0·5
β Leonis	11	41	54·98	54·98	0·00	74	38	43·96	43·9	+ 0·1
β Virginis	11	43	24·13	24·20	− 0·07	87	26	47·35	47·7	− 0·4
γ Ursæ Majoris	11	46	27·06	27·25	− 0·19	35	31	37·22	37·2	0·0
α Virginis	13	17	49·27	49·29	− 0·02	100	25	45·54	46·8	− 1·3
η Ursæ Majoris	13	42	1·27	1·27	0·00	39	59	12·81	13·4	− 0·6
α Boötis	14	9	16·62	16·55	+ 0·07	70	5	13·40	12·8	+ 0·6
8 Libræ	14	42	56·87	56·89	− 0·02	105	24	46·90	47·1	− 0·2
α Libræ	14	43	8·31	8·32	− 0·01	105	27	27·41	27·8	− 0·4
β Ursæ Minoris	14	51	9·16	8·89	+ 0·27	15	16	20·65	21·2	− 0·5
α Coronæ	15	28	45·69	45·67	+ 0·02	62	48	42·96	42·5	+ 0·5
α Serpentis	15	37	22·44	22·43	+ 0·01	83	7	52·88	53·0	− 0·1
α Scorpii	16	20	49·70	49·70	+ 0·00	116	7	2·96	2·9	+ 0·1
α^2 Herculis	17	8	16·24	16·23	+ 0·01	75	26	52·77	53·2	− 0·4
α Ophiuchi	17	28	26·22	26·13	+ 0·09	77	20	6·26	5·4	+ 0·9
γ Draconis	17	53	21·34	21·35	− 0·01	38	29	36·01	35·2	+ 0·8
α Lyræ	18	32	11·92	11·89	+ 0·03	51	20	40·73	39·6	+ 1·1
γ Aquilæ	19	39	36·22	36·13	+ 0·09	79	43	31·05	30·7	+ 0·4
α Aquilæ	19	43	57·12	57·08	+ 0·04	81	29	55·32	54·5	+ 0·8
β Aquilæ	19	48	26·14	26·19	− 0·05	83	56	24·92	25·1	− 0·2
α^1 Capricorni	20	9	53·12	53·16	− 0·04	102	56	16·17	18·7	− 2·5
α^2 Capricorni	20	10	17·03	17·08	− 0·05	102	58	33·07	35·2	− 2·1

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1870MNRAS..30..184P

	Greenwich Mean Right Ascension.			Rad.	G.—R.	Mean N.P.D.			Rad.	G.—R,
	h	m	s	″	″	°	′	″	″	″
α Cygni	20	36	39·60	39·47	− 0·13	45	13	6·23	7·3	− 1·1
α Cephei	21	15	14·13	14·18	− 0·05	28	0	24·59	24·8	− 0·2
α Aquarii	21	58	35·50	35·49	+ 0·01	90	59	54·48	55·1	− 0·6
α Piscis Australis	22	49	54·38	54·40	− 0·02	120	21	48·31	47·2	+ 1·1
α Pegasi	22	57	47·33	47·33	0·00	75	32	50·49	50·7	− 0·2
α Ursæ Minoris	1	8	3·05	2·37	+ 0·68	1	26	12·90	13·1	− 0·2
δ Ursæ Minoris	18	17	30·19	30·60	− 0·41	3	23	53·19	52·6	+ 0·6

The agreement, on the whole, between the Greenwich places and those of the *Tabulæ Reductionum* appears rather closer than with the *Radcliffe Observations*.

On the Resolvability of Star-Groups, regarded as a Test of Distance. By Richard A. Proctor, B.A.

There are considerations connected with the resolvability of star-groups which have not hitherto received much attention, so far as I am aware. They bear somewhat importantly on the opinions we are to form respecting the distribution of matter throughout the sidereal system.

In the first place, the resolvability of such clustering aggregations of stars as obviously form part of the sidereal system has been regarded as an important means of estimating the relative extension of different parts of that system. So long as a portion of such a clustering aggregation remains unresolved it has been assumed that the limits of the system in that direction lie beyond the range of the telescope which thus fails to effect resolution, and therefore that the extension of the system in that direction is far greater than in other directions where the same telescope shows the stars projected discretely on a perfectly black background.

In the second place, in the case of definite groups of stars, which either lie beyond the limits of the sidereal system, or if within those limits are yet separated from other parts of the system, and surrounded on all sides by relatively barren regions, it has been commonly assumed that we have, in the telescopic powers necessary to effect resolution, a means of forming a general estimate of the distances at which such groups may lie.

It is my purpose here to indicate certain considerations which point to opposite conclusions as respects both these cases.

If we were to accept the conclusion that where a portion of the galaxy is not resolvable with powerful telescopic aid, the sidereal system has a relatively great extension, it would follow from the smallness of the areas which many of these portions of the galaxy present that there is an extension of the system in those directions into long spike-shaped projections, lying in a di-

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